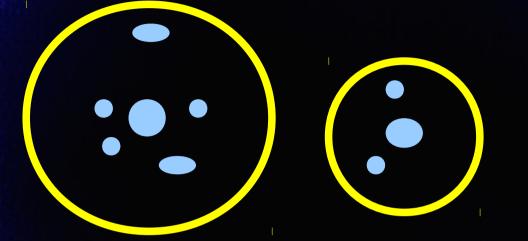
Star Formation and Quenching vs. Environment and Mass

Joanna Woo Avishai Dekel Sandra Faber David Koo et al.

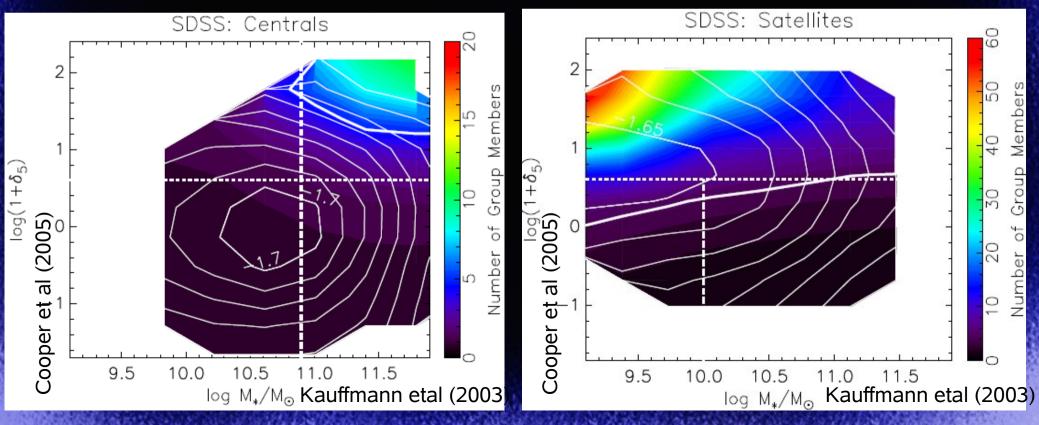
Goals

- Study relations between δ_N , M_* , M_h at $z\sim 0$, $z\sim 1$:
 - Splitting centrals and satellites
 - Using dual mode understanding of δ_{N}
 - Multi-halo mode, single-halo mode
- Study quenching vs. δ_N , M_* , M_h
 - Centrals vs. satellites, single halo/ multi halo mode
- Results:
 - $-M_h$ dominates quenching
 - Other kinds of quenching: small deviations

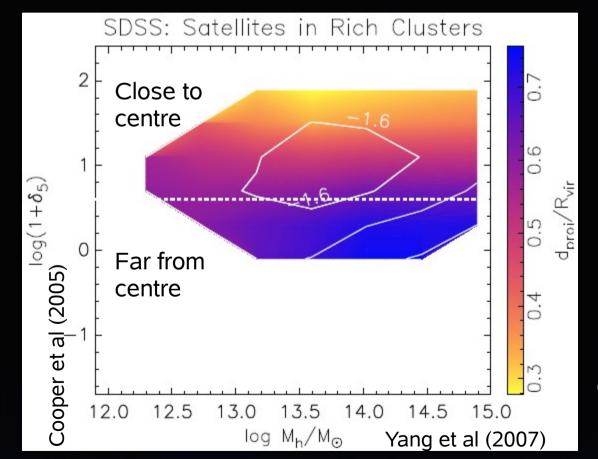
Two modes of δ_N



 δ_N mode depends on: -N (SDSS: N=5) -mag limit -z



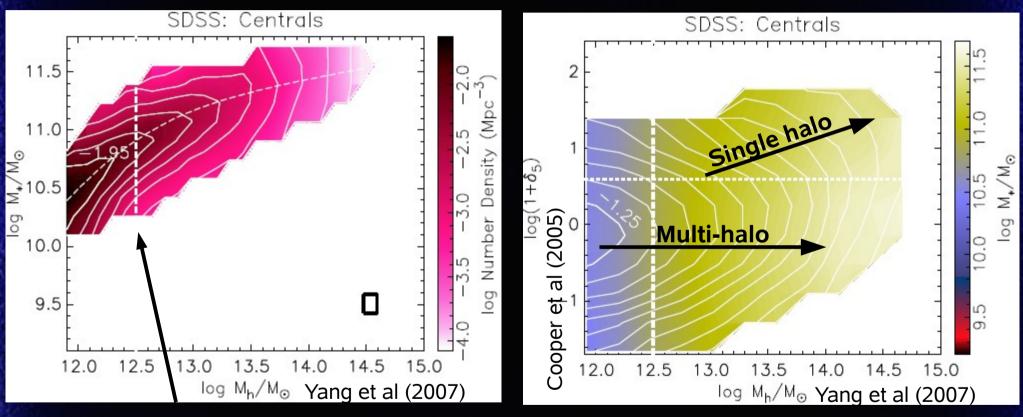
$\boldsymbol{\delta}_{N}$ also measures proximity to halo centre



(references for data labelled beside all axes)

(Only for groups with more than 5 members)

Relations between δ_N , M_* , M_h : Centrals

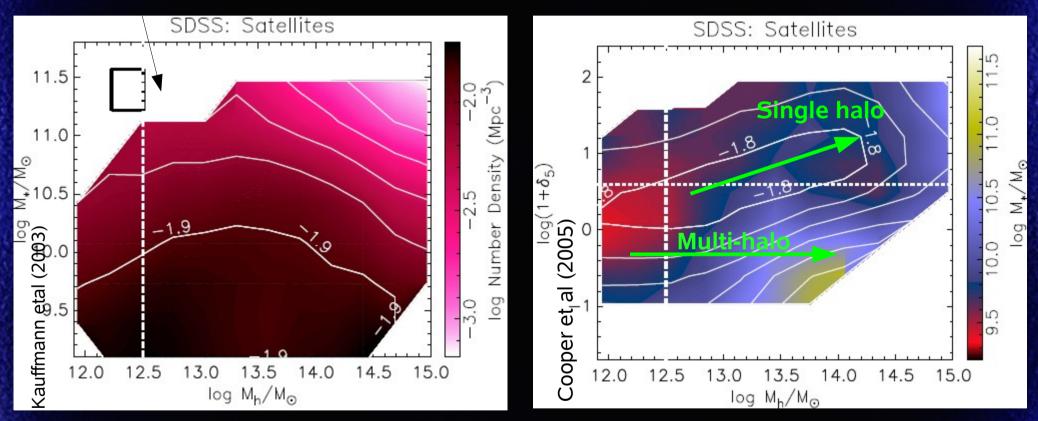


Shock scale ~ where relation begins to flatten

Expect δ_{N} to correlate with M_{h} within halo due to density profile

Relations between δ_N , M_* , M_h : Satellites

No big satellites in small halos



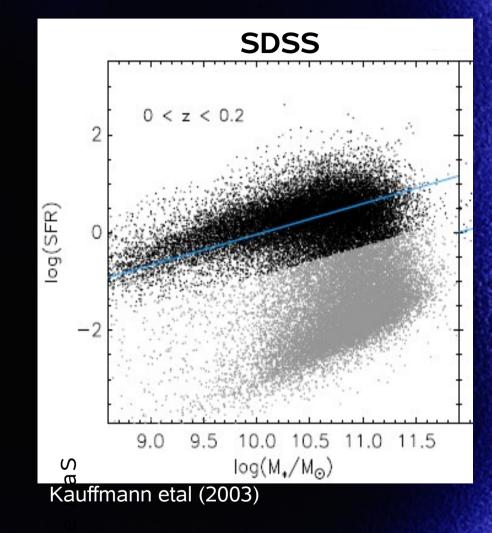
 $\delta_{_N}$ correlates with $M_{_h}$ also because of density profile

SFR Sequence

What's a "star-forming" galaxy? -on the SFR sequence

What's a "quenched" galaxy? - >2σ below SFR sequence

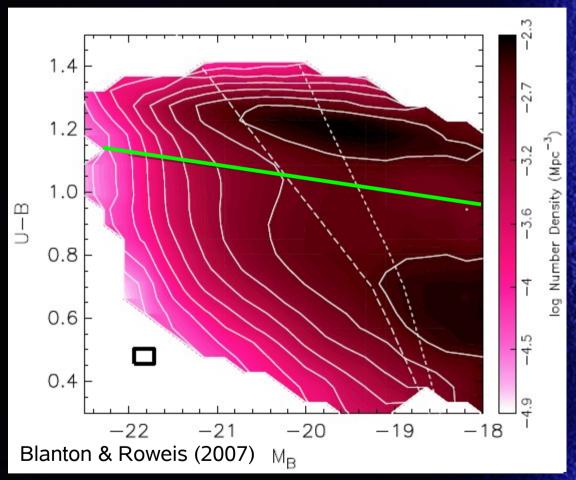
The "quenched fraction" = quenched / total (weighted by volume and spectroscopic completeness)



(references for data labelled beside all axes)

SFR Sequence vs. Blue Cloud

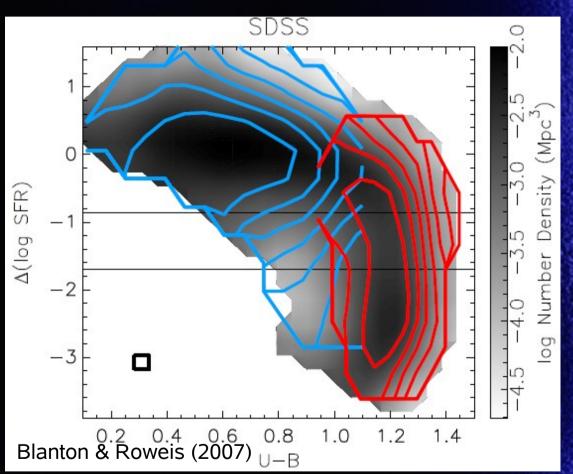
- How do the BC and RS compare to being on the SFR sequence or below it?
- 31% of the RS lie on the SFR sequence \rightarrow dusty!
- Only 7% of BC lie below SFR sequence (poststarburst)



The RS is not all quenched! Better to use SFR-M_{*} diagram to distinguish between "star-forming" and "quenched" galaxies.

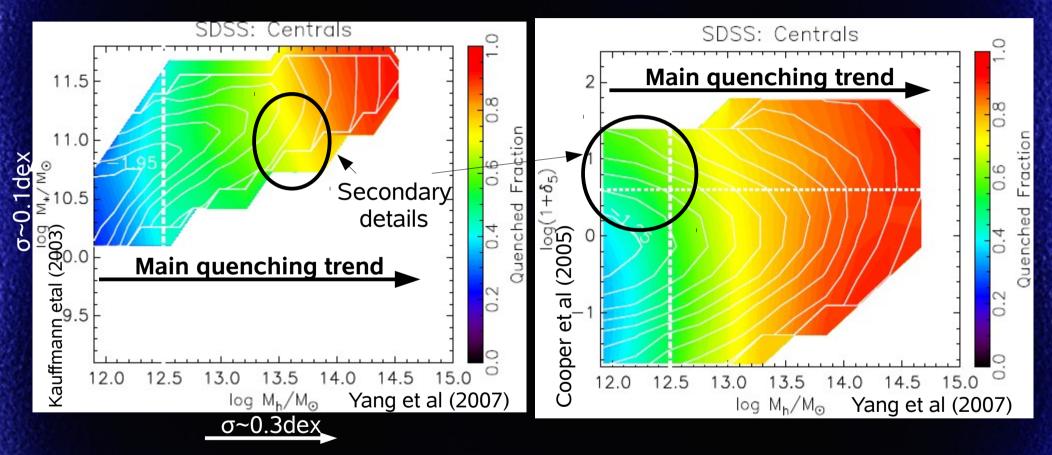
SFR Sequence vs. Blue Cloud

- How do the BC and RS compare to being on the SFR sequence or below it?
- 31% of the RS lie on the SFR sequence \rightarrow dusty!
- Only 7% of BC lie below SFR sequence (poststarburst)



The RS is not all quenched! Better to use SFR-M_{*} diagram to distinguish between "star-forming" and "quenched" galaxies.

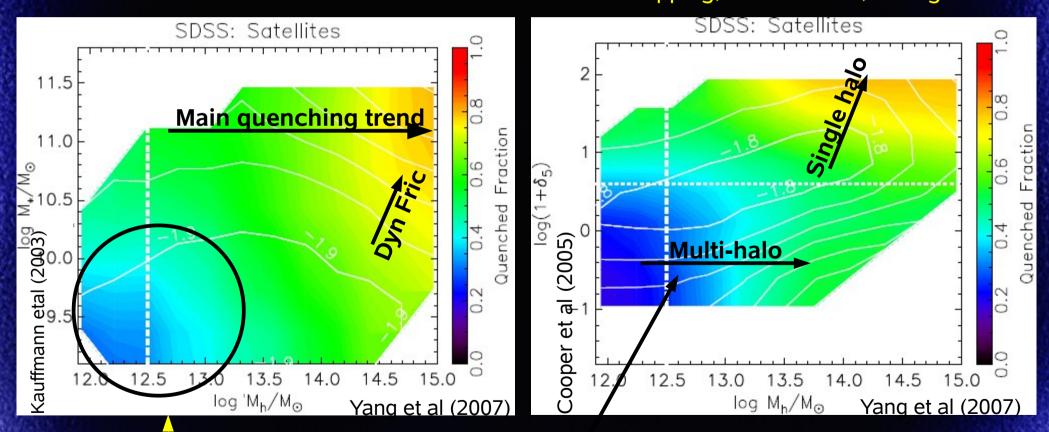
Quenching vs. δ_N , M_* , M_h : Centrals



The dominant trend of quenching is with *halo mass*, not stellar mass or environment density.

Quenching vs. δ_N , M_* , M_h : Satellites

Proximity to centre: -shock heating more efficient -stripping, harrassment, strangulation

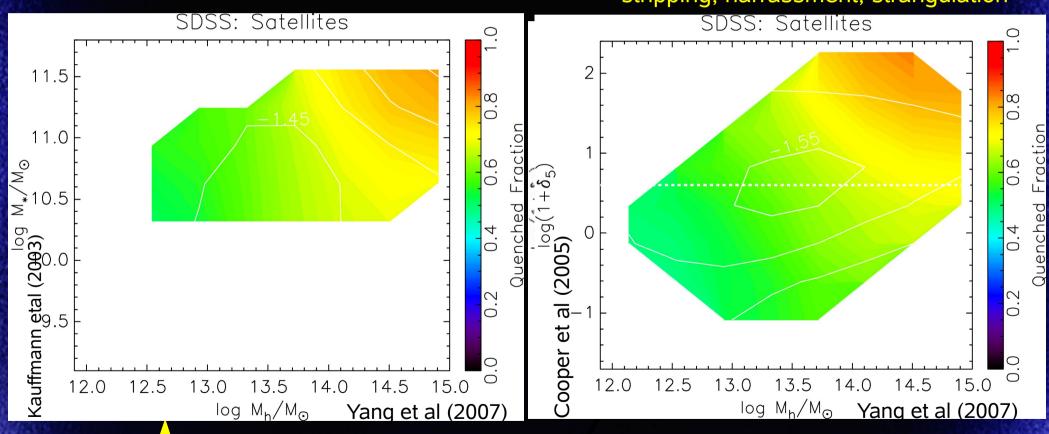


Here δ has nothing to do with the halo

Incompleteness?

Quenching vs. δ_N , M_* , M_h : Satellites

Proximity to centre: -shock heating more efficient -stripping, harrassment, strangulation



Here δ has nothing to do with the halo

Incompleteness?

Summary of Results

- The dominant quenching trend for centrals is with <u>halo mass</u> not M_* or δ_5
 - External quenching
 - Small dependence on M_* and δ_5 : deviation
- The quenching trend for satellites is with <u>**both**</u> M_h and δ_5 in the single halo mode and with M_h in the multi-halo mode
 - External quenching
 - Small M_{*} trend: consistent with dynamical friction
- Note: these quenching trends with M_h show up when measured by the quenched fraction instead of the red fraction;
 - Since large SF galaxies are dusty, the red fraction will correlate more with M_{*} than M_h causing quenching appear to be internal